

Black holes, the Van der Waals gas, compressibility and the speed of sound

Brian P. Dolan

Dept. of Mathematics, Heriot-Watt University, Edinburgh
and
Maxwell Institute for Mathematical Sciences, Edinburgh

Corfu, 14th Sept. 2013

Outline

Review of black hole thermodynamics

Basics

1st and 2nd laws

Smarr relation

Pressure and Enthalpy

Enthalpy and the 1st law

Volume

Equation of state

Critical behaviour

Compressibility

Speed of sound

Conclusions

Basics

- Entropy: $S \propto \frac{A}{\ell_{Pl}^2}$ ($\ell_{Pl}^2 = \hbar G/c^3$, $G = c = 1$).
Bekenstein (1972)
- Temperature, $T = \frac{\kappa \hbar}{2\pi}$: κ = surface gravity. Hawking (1974)

Schwarzschild black-hole, $\kappa = \frac{1}{4M}$

$$T = \frac{\hbar}{8\pi M}.$$

Basics

- Entropy: $S = \frac{1}{4} \frac{A}{\ell_{Pl}^2}$ ($\ell_{Pl}^2 = \hbar G/c^3$, $G = c = 1$).
Bekenstein (1972)
- Temperature, $T = \frac{\kappa \hbar}{2\pi}$: κ = surface gravity. Hawking (1974)

Schwarzschild black-hole, $\kappa = \frac{1}{4M}$

$$T = \frac{\hbar}{8\pi M}.$$

Basics

- Entropy: $S = \frac{1}{4} \frac{A}{\ell_{Pl}^2}$ ($\ell_{Pl}^2 = \hbar G/c^3$, $G = c = 1$).
Bekenstein (1972)
- Temperature, $T = \frac{\kappa \hbar}{2\pi}$: $\kappa = \text{surface gravity}$. Hawking (1974)

Schwarzschild black-hole, $\kappa = \frac{1}{4M}$

$$T = \frac{\hbar}{8\pi M}.$$

First and Second laws

- Internal energy $U(S)$, $T = \frac{\partial U}{\partial S}$:
identify $M = U(S) \Rightarrow dM = TdS$.
More generally $M = U(S, J, Q)$, (angular momentum,
electric charge),

First Law of Black Hole Thermodynamics

$$dM = T dS + \Omega dJ + \Phi dQ$$

First and Second laws

- Internal energy $U(S)$, $T = \frac{\partial U}{\partial S}$:
identify $M = U(S) \Rightarrow dM = TdS$.
More generally $M = U(S, J, Q)$, (angular momentum,
electric charge),

First Law of Black Hole Thermodynamics

$$dM = T dS + \Omega dJ + \Phi dQ$$

Smarr relation

Smarr (1973)

- Ordinary thermodynamics: $U(S, V, n)$ (n = number of moles) is a function of **extensive variables**. U is also extensive \Rightarrow

$$\lambda^d U(S, V, n) = U(\lambda^d S, \lambda^d V, \lambda^d n)$$

$$\Rightarrow U = S \frac{\partial U}{\partial S} + V \frac{\partial U}{\partial V} + n \frac{\partial U}{\partial n} \quad \text{Euler equation}$$

$$\Rightarrow U = ST - VP + n\mu \quad (\mu = \text{chemical potential})$$

$$\Rightarrow G = U + VP - ST = n\mu. \quad \text{Gibbs-Duhem relation}$$

Black hole in D dimensions, angular momenta J_i : $(D-2)M$

$$S \rightarrow \lambda^{D-2} S, J_i \rightarrow \lambda^{D-2} J_i, M \rightarrow \lambda^{D-3} M \Rightarrow$$

$$(D-3)M = (D-2)S \frac{\partial M}{\partial S} + (D-2)J_i \frac{\partial M}{\partial J_i}$$

$$(D-3)M = (D-2)S \frac{\partial M}{\partial S} + (D-2)J_i \frac{\partial M}{\partial J_i}$$

$$\Rightarrow (D-3)M = (D-2)ST + (D-2)J_i \Omega_i.$$

Smarr relation

Smarr (1973)

- Ordinary thermodynamics: $U(S, V, n)$ (n = number of moles) is a function of **extensive variables**. U is also extensive \Rightarrow

$$\lambda^d U(S, V, n) = U(\lambda^d S, \lambda^d V, \lambda^d n)$$

$$\Rightarrow U = S \frac{\partial U}{\partial S} + V \frac{\partial U}{\partial V} + n \frac{\partial U}{\partial n} \quad \text{Euler equation}$$

$$\Rightarrow U = ST - VP + n\mu \quad (\mu = \text{chemical potential})$$

$$\Rightarrow G = U + VP - ST = n\mu. \quad \text{Gibbs-Duhem relation}$$

Black hole in D dimensions, angular momenta J_i : $(D-3)M = (D-2)S + (D-2)J_i \Omega_i$

$$S \rightarrow \lambda^{D-2} S, J_i \rightarrow \lambda^{D-2} J_i, M \rightarrow \lambda^{D-3} M \Rightarrow$$

$$(D-3)M = (D-2)S \frac{\partial M}{\partial S} + (D-2)J_i \frac{\partial M}{\partial J_i}$$

$$\Rightarrow (D-3)M = (D-2)ST + (D-2)J_i \Omega_i.$$

Smarr relation

Smarr (1973)

- Ordinary thermodynamics: $U(S, V, n)$ (n = number of moles) is a function of **extensive variables**. U is also extensive \Rightarrow

$$\lambda^d U(S, V, n) = U(\lambda^d S, \lambda^d V, \lambda^d n)$$

$$\Rightarrow U = S \frac{\partial U}{\partial S} + V \frac{\partial U}{\partial V} + n \frac{\partial U}{\partial n} \quad \text{Euler equation}$$

$$\Rightarrow U = ST - VP + n\mu \quad (\mu = \text{chemical potential})$$

$$\Rightarrow G = U + VP - ST = n\mu. \quad \text{Gibbs-Duhem relation}$$

Black hole in D dimensions, angular momenta J_i : $U = U(S, V, J_i)$

$$S \rightarrow \lambda^{D-2} S, J_i \rightarrow \lambda^{D-2} J_i, M \rightarrow \lambda^{D-3} M \Rightarrow$$

$$(D-3)M = (D-2)S \frac{\partial U}{\partial S} + (D-2)J_i \frac{\partial U}{\partial J_i}$$

$$\Rightarrow (D-3)M = (D-2)ST + (D-2)J_i \Omega_i.$$

Smarr relation

Smarr (1973)

- Ordinary thermodynamics: $U(S, V, n)$ (n = number of moles) is a function of **extensive variables**. U is also extensive \Rightarrow

$$\lambda^d U(S, V, n) = U(\lambda^d S, \lambda^d V, \lambda^d n)$$

$$\Rightarrow U = S \frac{\partial U}{\partial S} + V \frac{\partial U}{\partial V} + n \frac{\partial U}{\partial n} \quad \text{Euler equation}$$

$$\Rightarrow U = ST - VP + n\mu \quad (\mu = \text{chemical potential})$$

$$\Rightarrow G = U + VP - ST = n\mu. \quad \text{Gibbs-Duhem relation}$$

Black hole in D dimensions, angular momenta J_i :

$$S \rightarrow \lambda^{D-2} S, J_i \rightarrow \lambda^{D-2} J_i, M \rightarrow \lambda^{D-3} M \Rightarrow$$

$$(D-3)M = (D-2)S \frac{\partial M}{\partial S} + (D-2)J_i \frac{\partial M}{\partial J_i}$$

$$\Rightarrow (D-3)M = (D-2)ST + (D-2)J_i \Omega_i.$$

Smarr relation

Smarr (1973)

- Ordinary thermodynamics: $U(S, V, n)$ (n = number of moles) is a function of **extensive variables**. U is also extensive \Rightarrow

$$\lambda^d U(S, V, n) = U(\lambda^d S, \lambda^d V, \lambda^d n)$$

$$\Rightarrow U = S \frac{\partial U}{\partial S} + V \frac{\partial U}{\partial V} + n \frac{\partial U}{\partial n} \quad \text{Euler equation}$$

$$\Rightarrow U = ST - VP + n\mu \quad (\mu = \text{chemical potential})$$

$$\Rightarrow G = U + VP - ST = n\mu. \quad \text{Gibbs-Duhem relation}$$

- Black hole in D dimensions, angular momenta J_i : ($Q = 0$)

$$S \rightarrow \lambda^{D-2} S, J_i \rightarrow \lambda^{D-2} J_i, M \rightarrow \lambda^{D-3} M \Rightarrow$$

$$\lambda^{D-3} M(S, J_i) = M(\lambda^{D-3} S, \lambda^{D-2} J_i)$$

$$(D-3)M = (D-2)S \frac{\partial M}{\partial S} + (D-2)J_i \frac{\partial M}{\partial J_i}$$

$$\Rightarrow (D-3)M = (D-2)ST + (D-2)J_i \Omega_i.$$

Smarr relation

Smarr (1973)

- Ordinary thermodynamics: $U(S, V, n)$ (n = number of moles) is a function of **extensive variables**. U is also extensive \Rightarrow

$$\lambda^d U(S, V, n) = U(\lambda^d S, \lambda^d V, \lambda^d n)$$

$$\Rightarrow U = S \frac{\partial U}{\partial S} + V \frac{\partial U}{\partial V} + n \frac{\partial U}{\partial n} \quad \text{Euler equation}$$

$$\Rightarrow U = ST - VP + n\mu \quad (\mu = \text{chemical potential})$$

$$\Rightarrow G = U + VP - ST = n\mu. \quad \text{Gibbs-Duhem relation}$$

- Black hole in D dimensions, angular momenta J_i : ($Q = 0$)

$$S \rightarrow \lambda^{D-2} S, J_i \rightarrow \lambda^{D-2} J_i, M \rightarrow \lambda^{D-3} M \Rightarrow$$

$$\lambda^{D-3} M(S, J_i) = M(\lambda^{D-3} S, \lambda^{D-2} J_i)$$

$$(D-3)M = (D-2)S \frac{\partial M}{\partial S} + (D-2)J_i \frac{\partial M}{\partial J_i}$$

$$\Rightarrow (D-3)M = (D-2)ST + (D-2)J_i \Omega_i.$$

Smarr relation

Smarr (1973)

- Ordinary thermodynamics: $U(S, V, n)$ (n = number of moles) is a function of **extensive variables**. U is also extensive \Rightarrow

$$\lambda^d U(S, V, n) = U(\lambda^d S, \lambda^d V, \lambda^d n)$$

$$\Rightarrow U = S \frac{\partial U}{\partial S} + V \frac{\partial U}{\partial V} + n \frac{\partial U}{\partial n} \quad \text{Euler equation}$$

$$\Rightarrow U = ST - VP + n\mu \quad (\mu = \text{chemical potential})$$

$$\Rightarrow G = U + VP - ST = n\mu. \quad \text{Gibbs-Duhem relation}$$

- Black hole in D dimensions, angular momenta J_i : ($Q = 0$)

$$S \rightarrow \lambda^{D-2} S, J_i \rightarrow \lambda^{D-2} J_i, M \rightarrow \lambda^{D-3} M \Rightarrow$$

$$\lambda^{D-3} M(S, J_i) = M(\lambda^{D-3} S, \lambda^{D-2} J_i)$$

$$(D-3)M = (D-2)S \frac{\partial M}{\partial S} + (D-2)J_i \frac{\partial M}{\partial J_i}$$

$$\Rightarrow (D-3)M = (D-2)ST + (D-2)J_i \Omega_i.$$

Smarr relation

- **Doesn't work when there's a Cosmological Constant!**
- Solution: (Kastor, Ray+Traschen [0904.2765])
include Λ as a thermodynamic variable, $\Lambda \rightarrow \lambda^{-2}\Lambda$,
(Henneaux+Teitelboim (1984))

$$\begin{aligned}\lambda^{D-3}M(S, \Lambda, J_i) &= M(\lambda^{D-2}S, \lambda^{-2}\Lambda, \lambda^{D-2}J_i) \\ (D-3)M &= (D-2)S \frac{\partial M}{\partial S} - 2\Lambda \frac{\partial M}{\partial \Lambda} + (D-2)J_i \frac{\partial M}{\partial J_i} \\ \Rightarrow (D-3)M &= (D-2)ST - 2\Lambda\Theta + (D-2)J_i \Omega_i.\end{aligned}$$

Smarr relation

- **Doesn't work when there's a Cosmological Constant!**
- Solution: (Kastor, Ray+Traschen [0904.2765])
include Λ as a thermodynamic variable, $\Lambda \rightarrow \lambda^{-2}\Lambda$,
(Henneaux+Teitelboim (1984))

$$\begin{aligned}\lambda^{D-3}M(S, \Lambda, J_i) &= M(\lambda^{D-2}S, \lambda^{-2}\Lambda, \lambda^{D-2}J_i) \\ (D-3)M &= (D-2)S \frac{\partial M}{\partial S} - 2\Lambda \frac{\partial M}{\partial \Lambda} + (D-2)J_i \frac{\partial M}{\partial J_i} \\ \Rightarrow (D-3)M &= (D-2)ST - 2\Lambda\Theta + (D-2)J_i \Omega_i.\end{aligned}$$

Smarr relation

- **Doesn't work when there's a Cosmological Constant!**
- Solution: (Kastor, Ray+Traschen [0904.2765])
include Λ as a thermodynamic variable, $\Lambda \rightarrow \lambda^{-2}\Lambda$,
(Henneaux+Teitelboim (1984))

$$\lambda^{D-3}M(S, \Lambda, J_i) = M(\lambda^{D-2}S, \lambda^{-2}\Lambda, \lambda^{D-2}J_i)$$

$$(D-3)M = (D-2)S \frac{\partial M}{\partial S} - 2\Lambda \frac{\partial M}{\partial \Lambda} + (D-2)J_i \frac{\partial M}{\partial J_i}$$

$$\Rightarrow (D-3)M = (D-2)ST - 2\Lambda\Theta + (D-2)J_i \Omega_i.$$

Smarr relation

- **Doesn't work when there's a Cosmological Constant!**
- Solution: (Kastor, Ray+Traschen [0904.2765])
include Λ as a thermodynamic variable, $\Lambda \rightarrow \lambda^{-2}\Lambda$,
(Henneaux+Teitelboim (1984))

$$\begin{aligned}\lambda^{D-3}M(S, \Lambda, J_i) &= M(\lambda^{D-2}S, \lambda^{-2}\Lambda, \lambda^{D-2}J_i) \\ (D-3)M &= (D-2)S \frac{\partial M}{\partial S} - 2\Lambda \frac{\partial M}{\partial \Lambda} + (D-2)J_i \frac{\partial M}{\partial J_i} \\ \Rightarrow (D-3)M &= (D-2)ST - 2\Lambda\Theta + (D-2)J_i \Omega_i.\end{aligned}$$

Smarr relation

- **Doesn't work when there's a Cosmological Constant!**
- **Solution:** (Kastor, Ray+Traschen [0904.2765])
include Λ as a thermodynamic variable, $\Lambda \rightarrow \lambda^{-2}\Lambda$,
(Henneaux+Teitelboim (1984))

$$\begin{aligned}\lambda^{D-3}M(S, \Lambda, J_i) &= M(\lambda^{D-2}S, \lambda^{-2}\Lambda, \lambda^{D-2}J_i) \\ (D-3)M &= (D-2)S \frac{\partial M}{\partial S} - 2\Lambda \frac{\partial M}{\partial \Lambda} + (D-2)J_i \frac{\partial M}{\partial J_i} \\ \Rightarrow (D-3)M &= (D-2)ST - 2\Lambda\Theta + (D-2)J_i \Omega_i.\end{aligned}$$

$$\Theta := \left. \frac{\partial M}{\partial \Lambda} \right|_{S, J}$$

Enthalpy

- Pressure $P = -\frac{\Lambda}{8\pi} \Rightarrow$ mass is the **enthalpy**:

Enthalpy

$$M = H(S, P, J)$$

(Kastor, Ray+Traschen [0904.2765])

- Internal energy $U(S, V, J)$ is the Legendre transform:

$$U = H - PV$$

where $V := \frac{\partial M}{\partial P}$ is a thermodynamic volume.

First Law of thermodynamics

$$dU = T dS - P dV + \Omega dJ$$

- For $\Lambda < 0$, Penrose process can be up to 50% efficient (0.29% for $\Lambda = 0$). (BPD [1008.5023])

Enthalpy

- Pressure $P = -\frac{\Lambda}{8\pi} \Rightarrow$ mass is the **enthalpy**:

Enthalpy

$$M = H(S, P, J)$$

(Kastor, Ray+Traschen [0904.2765])

- Internal energy $U(S, V, J)$ is the Legendre transform:

$$U = H - PV$$

where $V := \frac{\partial M}{\partial P}$ is a thermodynamic volume.

First Law of thermodynamics

$$dU = T dS - P dV + \Omega dJ$$

- For $\Lambda < 0$, Penrose process can be up to 50% efficient (0.29% for $\Lambda = 0$). (BPD [1008.5023])

Enthalpy

- Pressure $P = -\frac{\Lambda}{8\pi} \Rightarrow$ mass is the **enthalpy**:

Enthalpy

$$M = H(S, P, J)$$

(Kastor, Ray+Traschen [0904.2765])

- Internal energy $U(S, V, J)$ is the Legendre transform:

$$U = H - PV$$

where $V := \frac{\partial M}{\partial P}$ is a **thermodynamic volume**.

First Law of thermodynamics

$$dU = T dS - P dV + \Omega dJ$$

- For $\Lambda < 0$, Penrose process can be up to 50% efficient (0.29% for $\Lambda = 0$). (BPD [1008.5023])

Enthalpy

- Pressure $P = -\frac{\Lambda}{8\pi} \Rightarrow$ mass is the **enthalpy**:

Enthalpy

$$M = H(S, P, J)$$

(Kastor, Ray+Traschen [0904.2765])

- Internal energy $U(S, V, J)$ is the Legendre transform:

$$U = H - PV$$

where $V := \frac{\partial M}{\partial P}$ is a **thermodynamic volume**.

First Law of thermodynamics

$$dU = T dS - P dV + \Omega dJ$$

- For $\Lambda < 0$, Penrose process can be up to 50% efficient (0.29% for $\Lambda = 0$). (BPD [1008.5023])

Enthalpy

- Pressure $P = -\frac{\Lambda}{8\pi} \Rightarrow$ mass is the **enthalpy**:

Enthalpy

$$M = H(S, P, J)$$

(Kastor, Ray+Traschen [0904.2765])

- Internal energy $U(S, V, J)$ is the Legendre transform:

$$U = H - PV$$

where $V := \frac{\partial M}{\partial P}$ is a **thermodynamic volume**.

First Law of thermodynamics

$$dU = T dS - P dV + \Omega dJ$$

- For $\Lambda < 0$, Penrose process can be up to 50% efficient (0.29% for $\Lambda = 0$). (BPD [1008.5023])

Thermodynamic Volume

- In $4 - D$ the thermodynamic volume evaluates to

$$V = \frac{1}{3} \left(r_h \mathcal{A}_h + \frac{J^2}{M} \right),$$

where \mathcal{A}_h is area of the event horizon and r_h the radius.
(Cvetič *et al.* [1012.2888].)

Equation of state $(D = 4, \Lambda < 0)$

- Critical point ($J \neq 0$):

$$(PJ)_{crit} \approx 0.002857, \quad (T\sqrt{J})_{crit} \approx 0.04175, \quad \left(\frac{V}{J^{3/2}}\right)_{crit} \approx 115.8$$

- Define

$$t := \frac{T - T_c}{T_c}, \quad v := \frac{V - V_c}{V_c}, \quad p := \frac{P - P_c}{P_c}.$$

Expand the equation of state about the critical point:

$$p = 2.42t - 0.81tv - 0.21v^3 + o(t^2, tv^2, v^4).$$

cf. Van der Waals gas: $p = 4t - 6tv - \frac{3}{2}v^3 + o(t^2, tv^2, v^4)$.

(Gunasekaran *et al.* [1208.6251], BPD [1209.1272]).

Compressibility

- Adiabatic compressibility: $\kappa = -\frac{1}{V} \left. \frac{\partial V}{\partial P} \right|_{S, J}$
- Rotating black-hole in D -dimensions (Myers-Perry).
Dimensionless angular momenta, $\mathcal{J}_i := \frac{2\pi J_i}{S}$,
Constraint: $T \geq 0 \Rightarrow \sum_i \frac{1}{1+\mathcal{J}_i^2} \geq \begin{cases} \frac{1}{2} & \text{even } D \\ 1 & \text{odd } D. \end{cases}$

Compressibility, $\Lambda \rightarrow 0$

$$\kappa = \frac{16\pi r_h^2}{(D-1)(D-2)^2} \left\{ \frac{(D-2) \sum_i \mathcal{J}_i^4 - (\sum_i \mathcal{J}_i^2)^2}{D-2 + \sum_i \mathcal{J}_i^2} \right\},$$

(BPD [1308.5403])

- $0 \leq \kappa < \infty$.

Compressibility

- Adiabatic compressibility: $\kappa = -\frac{1}{V} \left. \frac{\partial V}{\partial P} \right|_{S, J}$
- Rotating black-hole in D -dimensions (Myers-Perry).

Dimensionless angular momenta, $\mathcal{J}_i := \frac{2\pi J_i}{S}$,

$$\text{Constraint: } T \geq 0 \Rightarrow \sum_i \frac{1}{1+\mathcal{J}_i^2} \geq \begin{cases} \frac{1}{2} & \text{even } D \\ 1 & \text{odd } D. \end{cases}$$

Compressibility, $\Lambda \rightarrow 0$

$$\kappa = \frac{16\pi r_h^2}{(D-1)(D-2)^2} \left\{ \frac{(D-2) \sum_i \mathcal{J}_i^4 - (\sum_i \mathcal{J}_i^2)^2}{D-2 + \sum_i \mathcal{J}_i^2} \right\},$$

(BPD [1308.5403])

- $0 \leq \kappa < \infty$.

Compressibility

- Adiabatic compressibility: $\kappa = -\frac{1}{V} \left. \frac{\partial V}{\partial P} \right|_{S, J}$
- Rotating black-hole in D -dimensions (Myers-Perry).
Dimensionless angular momenta, $\mathcal{J}_i := \frac{2\pi J_i}{S}$,

$$\text{Constraint: } T \geq 0 \Rightarrow \sum_i \frac{1}{1+\mathcal{J}_i^2} \geq \begin{cases} \frac{1}{2} & \text{even } D \\ 1 & \text{odd } D. \end{cases}$$

Compressibility, $\Lambda \rightarrow 0$

$$\kappa = \frac{16\pi r_h^2}{(D-1)(D-2)^2} \left\{ \frac{(D-2) \sum_i \mathcal{J}_i^4 - (\sum_i \mathcal{J}_i^2)^2}{D-2 + \sum_i \mathcal{J}_i^2} \right\},$$

(BPD [1308.5403])

- $0 \leq \kappa < \infty$.

Compressibility

- Adiabatic compressibility: $\kappa = -\frac{1}{V} \left. \frac{\partial V}{\partial P} \right|_{S, J}$
- Rotating black-hole in D -dimensions (Myers-Perry).
Dimensionless angular momenta, $\mathcal{J}_i := \frac{2\pi J_i}{S}$,

$$\text{Constraint: } T \geq 0 \Rightarrow \sum_i \frac{1}{1+\mathcal{J}_i^2} \geq \begin{cases} \frac{1}{2} & \text{even } D \\ 1 & \text{odd } D. \end{cases}$$

Compressibility, $\Lambda \rightarrow 0$

$$\kappa = \frac{16\pi r_h^2}{(D-1)(D-2)^2} \left\{ \frac{(D-2) \sum_i \mathcal{J}_i^4 - (\sum_i \mathcal{J}_i^2)^2}{D-2 + \sum_i \mathcal{J}_i^2} \right\},$$

(BPD [1308.5403])

- $0 \leq \kappa < \infty$.

Compressibility

- Adiabatic compressibility: $\kappa = -\frac{1}{V} \left. \frac{\partial V}{\partial P} \right|_{S, J}$
- Rotating black-hole in D -dimensions (Myers-Perry).
Dimensionless angular momenta, $\mathcal{J}_i := \frac{2\pi J_i}{S}$,

$$\text{Constraint: } T \geq 0 \Rightarrow \sum_i \frac{1}{1+\mathcal{J}_i^2} \geq \begin{cases} \frac{1}{2} & \text{even } D \\ 1 & \text{odd } D. \end{cases}$$

Compressibility, $\Lambda \rightarrow 0$

$$\kappa = \frac{16\pi r_h^2}{(D-1)(D-2)^2} \left\{ \frac{(D-2) \sum_i \mathcal{J}_i^4 - (\sum_i \mathcal{J}_i^2)^2}{D-2 + \sum_i \mathcal{J}_i^2} \right\},$$

(BPD [1308.5403])

- $0 \leq \kappa < \infty$.

Compressibility

- Adiabatic compressibility: $\kappa = -\frac{1}{V} \left. \frac{\partial V}{\partial P} \right|_{S, J}$
- Rotating black-hole in D -dimensions (Myers-Perry).
Dimensionless angular momenta, $\mathcal{J}_i := \frac{2\pi J_i}{S}$,

$$\text{Constraint: } T \geq 0 \Rightarrow \sum_i \frac{1}{1+\mathcal{J}_i^2} \geq \begin{cases} \frac{1}{2} & \text{even } D \\ 1 & \text{odd } D. \end{cases}$$

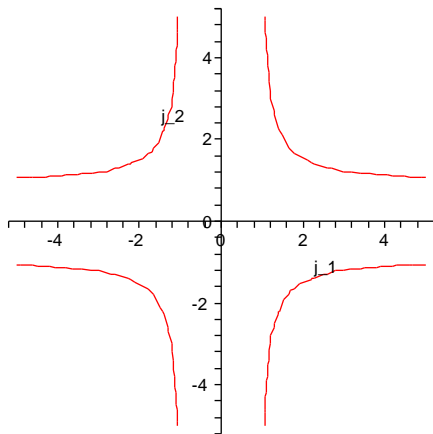
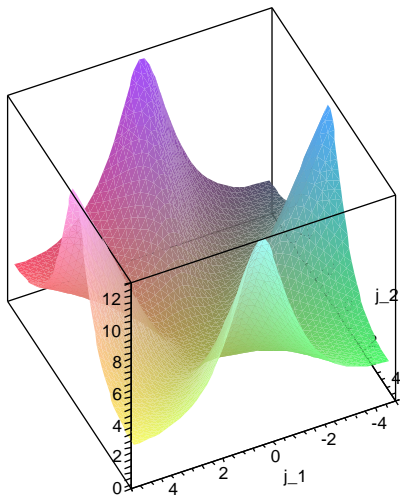
Compressibility, $\Lambda \rightarrow 0$

$$\kappa = \frac{16\pi r_h^2}{(D-1)(D-2)^2} \left\{ \frac{(D-2) \sum_i \mathcal{J}_i^4 - (\sum_i \mathcal{J}_i^2)^2}{D-2 + \sum_i \mathcal{J}_i^2} \right\},$$

(BPD [1308.5403])

- $0 \leq \kappa < \infty$.

Example of compressibility, $D = 6$, $SO(5)$: J_1, J_2 ,



Speed of sound

- Define $\rho := \frac{M}{V}$, then the thermodynamic speed of sound is

$$c_s^2 = \left. \frac{\partial P}{\partial \rho} \right|_{S,J}, \quad \Rightarrow \quad c_s^{-2} = 1 + \kappa \rho.$$

Speed of sound

$$c_s^2 = \frac{1}{(D-2)} \frac{(D-2 + \sum_i \mathcal{J}_i^2)^2}{(D-2 + 2 \sum_i \mathcal{J}_i^2 + \sum_i \mathcal{J}_i^4)}.$$

Speed of sound

- Define $\rho := \frac{M}{V}$, then the thermodynamic speed of sound is

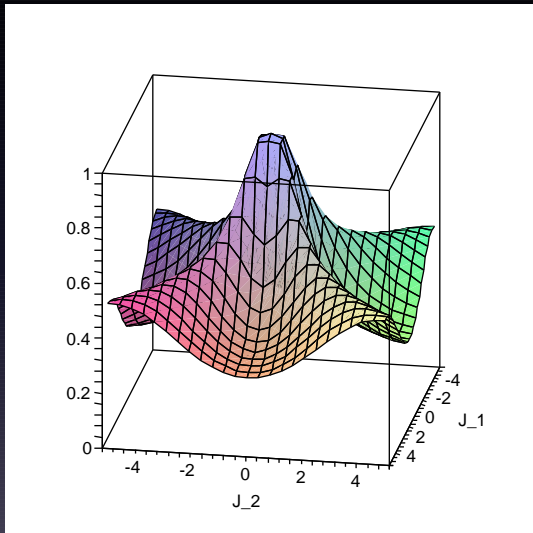
$$c_s^2 = \left. \frac{\partial P}{\partial \rho} \right|_{S,J}, \quad \Rightarrow \quad c_s^{-2} = 1 + \kappa \rho.$$

Speed of sound

$$c_s^2 = \frac{1}{(D-2)} \frac{(D-2 + \sum_i \mathcal{J}_i^2)^2}{(D-2 + 2\sum_i \mathcal{J}_i^2 + \sum_i \mathcal{J}_i^4)}.$$

- $\frac{1}{D-2} \leq c_s^2 \leq 1.$

c_s^2 in 6-dimensions:



Conclusions

- $\Lambda \neq 0 \Rightarrow P dV$ term in black hole 1st law.
- Black hole mass is identified with **enthalpy**, $H(S, P, J_i)$:

$$\begin{aligned}dM &= dH = T dS + V dP + \Omega_i J_i, \\dU &= T dS - P dV + \Omega_i J_i.\end{aligned}$$

- Hawking temperature: $T = \left(\frac{\partial H}{\partial S}\right)_P$.
- “Thermodynamic” volume: $V = \left(\frac{\partial H}{\partial P}\right)_T$.
- PdV term affects Penrose processes — more efficient in asymptotically AdS space-times.
- $D = 4$: **Van der Waals type equation of state**.
- Compressibility, $0 \leq \kappa < \infty$, with $\kappa \rightarrow \infty$ for some $J_i \rightarrow \infty$.
- Instability of ultra-spinning black-holes.
- Speed of sound: $\frac{1}{D-2} \leq c_s^2 \leq 1$.

Non-zero charge, zero J

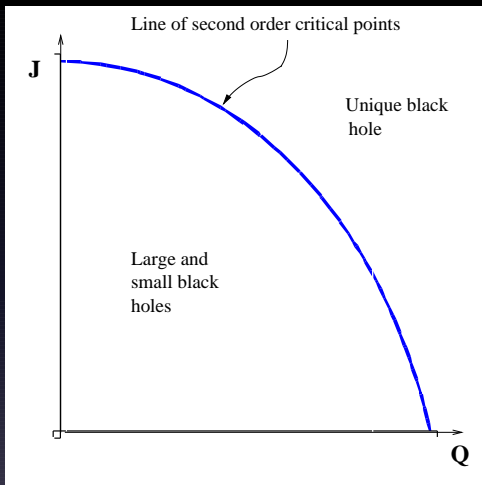
Reissner-Nordström — Anti-de Sitter ($J = 0$, $Q \neq 0$):

- Gibbs free energy is a ‘swallowtail’, same as Van der Waals
Champlin, Emparan, Johnson+Myers:
[hep-th/9902170]; [hep-th9904197].
- Equation of state:

$$p = \frac{8}{3}t - \frac{8}{9}tv - \frac{4}{81}v^3 + o(t^2, tv^2, v^4),$$

Critical exponents are mean field
(Kubizňák+Mann [arXiv:1205.0559]).

Kerr-Reissner-Nordström-AdS



Reissner-Nordström anti-de Sitter ($J \neq 0$, $Q \neq 0$).

No Hawking-Page phase transition in canonical ensemble, if $J \neq 0$ or $Q \neq 0$

Caldarelli, Gognola+Klemm [hep-th/9908022].